

Cite as: R. Orosei *et al.*, *Science*
10.1126/science.aar7268 (2018).

Radar evidence of subglacial liquid water on Mars

R. Orosei^{1*}, S. E. Lauro², E. Pettinelli², A. Cicchetti³, M. Coradini⁴, B. Cosciotti², F. Di Paolo¹, E. Flamini⁴, E. Mattei², M. Pajola⁵, F. Soldovieri⁶, M. Cartacci³, F. Cassenti⁷, A. Frigeri³, S. Giuppi³, R. Martuffi⁷, A. Masdea⁸, G. Mitri⁹, C. Nenna¹⁰, R. Noschese³, M. Restano¹¹, R. Seu⁷

¹Istituto di Radioastronomia, Istituto Nazionale di Astrofisica, Via Piero Gobetti 101, 40129 Bologna, Italy. ²Dipartimento di Matematica e Fisica, Università degli Studi Roma Tre, Via della Vasca Navale 84, 00146 Roma, Italy. ³Istituto di Astrofisica e Planetologia Spaziali, Istituto Nazionale di Astrofisica, Via del Fosso del Cavaliere 100, 00133 Roma, Italy. ⁴Agenzia Spaziale Italiana, Via del Politecnico, 00133 Roma, Italy. ⁵Osservatorio Astronomico di Padova, Istituto Nazionale di Astrofisica, Vicolo Osservatorio 5, 35122 Padova, Italy. ⁶Consiglio Nazionale delle Ricerche, Istituto per il Rilevamento Elettromagnetico dell'Ambiente, Via Diocleziano 328, 80124 Napoli, Italy. ⁷Dipartimento di Ingegneria dell'Informazione, Elettronica e Telecomunicazioni, Università degli Studi di Roma "La Sapienza," Via Eudossiana 18, 00184 Roma, Italy. ⁸E.P. Elettronica Progetti, Via Traspontina 25, 00040 Ariccia (RM), Italy. ⁹International Research School of Planetary Sciences, Università degli Studi "Gabriele d'Annunzio," Viale Pindaro 42, 65127 Pescara (PE), Italy. ¹⁰Danfoss Drives, Romstrasse 2 – Via Roma 2, 39014 Burgstall – Postal (BZ), Italy. ¹¹Serco, c/o ESA Centre for Earth Observation, Largo Galileo Galilei 1, 00044 Frascati (RM), Italy.

*Corresponding author. Email: roberto.oroisei@inaf.it

The presence of liquid water at the base of the martian polar caps has long been suspected but not observed. We surveyed the Planum Australe region using the MARSIS (Mars Advanced Radar for Subsurface and Ionosphere Sounding) instrument, a low-frequency radar on the Mars Express spacecraft. Radar profiles collected between May 2012 and December 2015 contain evidence of liquid water trapped below the ice of the South Polar Layered Deposits. Anomalously bright subsurface reflections are evident within a well-defined, 20-kilometer-wide zone centered at 193°E, 81°S, which is surrounded by much less reflective areas. Quantitative analysis of the radar signals shows that this bright feature has high relative dielectric permittivity (>15), matching that of water-bearing materials. We interpret this feature as a stable body of liquid water on Mars.

The presence of liquid water at the base of the martian polar caps was first hypothesized more than 30 years ago (1) and has been inconclusively debated ever since. Radio echo sounding (RES) is a suitable technique to resolve this dispute, because low-frequency radars have been used extensively and successfully to detect liquid water at the bottom of terrestrial polar ice sheets. An interface between ice and water, or alternatively between ice and water-saturated sediments, produces bright radar reflections (2, 3). The Mars Advanced Radar for Subsurface and Ionosphere Sounding (MARSIS) instrument on the Mars Express spacecraft (4) is used to perform RES experiments (5). MARSIS has surveyed the martian subsurface for more than 12 years in search of evidence of liquid water (6). Strong basal echoes have been reported in an area close to the thickest part of the South Polar Layered Deposits (SPLD), Mars' southern ice cap (7). These features were interpreted as due to the propagation of the radar signals through a very cold layer of pure water ice having negligible attenuation (7). Anomalously bright reflections were subsequently detected in other areas of the SPLD (8).

On Earth, the interpretation of radar data collected above the polar ice sheets is usually based on the combination of qualitative (the morphology of the bedrock) and quantitative (the reflected radar peak power) analyses (3, 9). The MARSIS design, particularly the very large footprint (~3 to 5 km), does not provide high spatial resolution, strongly limiting its ability to discriminate the presence of subglacial water bodies

from the shape of the basal topography (10). Therefore, an unambiguous detection of liquid water at the base of the polar deposit requires a quantitative estimation of the relative dielectric permittivity (hereafter, permittivity) of the basal material, which determines the radar echo strength.

Between 29 May 2012 and 27 December 2015, MARSIS surveyed a 200-km-wide area of Planum Australe, centered at 193°E, 81°S (Fig. 1), which roughly corresponds to a previous study area (8). This area does not exhibit any peculiar characteristics, either in topographic data from the Mars Orbiter Laser Altimeter (MOLA) (Fig. 1A) (11, 12) or in the available orbital imagery (Fig. 1B) (13). It is topographically flat, composed of water ice with 10 to 20% admixed dust (14, 15), and seasonally covered by a very thin layer of CO₂ ice that does not exceed 1 m in thickness (16, 17). In the same location, higher-frequency radar observations performed by the Shallow Radar instrument on the Mars Reconnaissance Orbiter (18), revealed barely any internal layering in the SPLD and did not detect any basal echo (fig. S1), in marked contrast with findings for the North Polar Layer Deposits and other regions of the SPLD (19).

A total of 29 radar profiles were acquired using the onboard unprocessed data mode (5) by transmitting closely spaced radio pulses centered at either 3 and 4 MHz or 4 and 5 MHz (table S1). Observations were performed when the spacecraft was on the night side of Mars to minimize ionospheric dispersion of the signal. Figure 2A shows an example

of a MARSIS radargram collected in the area, where the sharp surface reflection is followed by several secondary reflections produced by the interfaces between layers within the SPLD. The last of these echoes represents the reflection between the ice-rich SPLD and the underlying material (hereafter, basal material). In most of the investigated area, the basal reflection is weak and diffuse, but in some locations, it is very sharp and has a greater intensity (bright reflections) than the surrounding areas and the surface (Fig. 2B). Where the observations from multiple orbits overlap, the data acquired at the same frequency have consistent values of both surface and subsurface echo power (fig. S2).

The two-way pulse travel time between the surface and basal echoes can be used to estimate the depth of the subsurface reflector and map the basal topography. Assuming an average signal velocity of 170 m/ μ s within the SPLD, close to that of water ice (20), the depth of the basal reflector is about 1.5 km below the surface. The large size of the MARSIS footprint and the diffuse nature of basal echoes outside the bright reflectors prevent a detailed reconstruction of the basal topography, but a regional slope from west to east is recognizable (Fig. 3A). The subsurface area where the bright reflections are concentrated is topographically flat and surrounded by higher ground, except on its eastern side, where there is a depression.

The permittivity, which provides constraints on the composition of the basal material, can in principle be retrieved from the power of the reflected signal at the base of the SPLD. Unfortunately, the radiated power of the MARSIS antenna is unknown because it could not be calibrated on the ground (owing to the instrument's large dimensions), and thus the intensity of the reflected echoes can only be considered in terms of relative quantities. It is common to normalize the intensity of the subsurface echo to the surface value (21)—i.e., to compute the ratio between basal and surface echo power. Such a procedure has the advantage of also compensating for any ionospheric attenuation of the signal. Following this approach, we normalized the subsurface echo power to the median of the surface power computed along each orbit; we found that all normalized profiles at a given frequency yield consistent values of the basal echo power (fig. S3). Figure 3B shows a regional map of basal echo power after normalization; bright reflections are localized around 193°E, 81°S in all intersecting orbits, outlining a well-defined, 20-km-wide subsurface anomaly.

To compute the basal permittivity, we also require information about the dielectric properties of the SPLD, which depend on the composition and temperature of the deposits. Because the exact ratio between water ice and dust is unknown (15), and because the thermal gradient between the surface and the base of the SPLD is poorly constrained (22), we explored the range of plausible values for such parameters

and computed the corresponding range of permittivity values. The following general assumptions were made: (i) The SPLD is composed of a mixture of water ice and dust in varying proportions (from 2 to 20%), and (ii) the temperature profile inside the SPLD is linear, starting from a fixed temperature at the surface (160 K) and rising to a variable temperature at the base of the SPLD (range, 170 to 270 K). Various electromagnetic scenarios were computed (5) by considering a plane wave impinging normally onto a structure with three layers: a semi-infinite layer with the permittivity of free space, a homogeneous layer representing the SPLD, and another semi-infinite layer representing the material beneath the SPLD, with variable permittivity values. The output of this computation is an envelope encompassing a family of curves that relate the normalized basal echo power to the permittivity of the basal material (Fig. 4A). This envelope is used to determine the distribution of the basal permittivity (inside and outside the bright area) by weighting each admissible value of the permittivity with the values of the probability distribution of the normalized basal echo power (Fig. 4B). This procedure generated two distinct distributions of the basal permittivity estimated inside and outside the bright reflection area (Fig. 4C and fig. S4), whose median values at 3, 4, and 5 MHz are 30 ± 3 , 33 ± 1 , and 22 ± 1 and 9.9 ± 0.5 , 7.5 ± 0.1 , and 6.7 ± 0.1 , respectively. The basal permittivity outside the bright area is in the range of 4 to 15, typical for dry terrestrial volcanic rocks. It is also in agreement with previous estimates of 7.5 to 8.5 for the material at the base of the SPLD (23) and with values derived from radar surface echo power for dense dry igneous rocks on the martian surface at midlatitudes (24, 25). Conversely, permittivity values as high as those found within the bright area have not previously been observed on Mars. On Earth, values greater than 15 are seldom associated with dry materials (26). RES data collected in Antarctica (27) and Greenland (9) show that a permittivity larger than 15 is indicative of the presence of liquid water below polar deposits. On the basis of the evident analogy of the physical phenomena on Earth and Mars, we can infer that the high permittivity values retrieved for the bright area below the SPLD are due to (partially) water-saturated materials and/or layers of liquid water.

We examined other possible explanations for the bright area below the SPLD (supplementary text). For example, a CO₂ ice layer at the top or the bottom of the SPLD, or a very low temperature of the H₂O ice throughout the SPLD, could enhance basal echo power compared with surface reflections. We reject these explanations (supplementary text), either because of the very specific and unlikely physical conditions required, or because they do not cause sufficiently strong basal reflections (figs. S5 and S6). Although the pressure and the temperature at the base of the SPLD would be compatible

with the presence of liquid CO₂, its relative dielectric permittivity is much lower (about 1.6) (28) than that of liquid water (about 80), so it does not produce bright reflections.

The substantial amounts of magnesium, calcium, and sodium perchlorate in the soil of the northern plains of Mars, discovered using the Phoenix lander's Wet Chemistry Lab (29), support the presence of liquid water at the base of the polar deposits. Perchlorates can form through different physical and/or chemical mechanisms (30, 31) and have been detected in different areas of Mars. It is therefore reasonable to assume that they are also present at the base of the SPLD. Because the temperature at the base of the polar deposits is estimated to be around 205 K (32), and because perchlorates strongly suppress the freezing point of water (to a minimum of 204 and 198 K for magnesium and calcium perchlorates, respectively) (29), we therefore find it plausible that a layer of perchlorate brine could be present at the base of the polar deposits. The brine could be mixed with basal soils to form a sludge or could lie on top of the basal material to form localized brine pools (32).

The lack of previous radar detections of subglacial liquid water has been used to support the hypothesis that the polar caps are too thin for basal melting and has led some authors to state that liquid water may be located deeper than previously thought [e.g., (33)]. The MARSIS data show that liquid water can be stable below the SPLD at relatively shallow depths (about 1.5 km), thus constraining models of Mars' hydrosphere. The limited raw-data coverage of the SPLD (a few percent of the area of Planum Australe) and the large size required for a meltwater patch to be detectable by MARSIS (several kilometers in diameter and several tens of centimeters in thickness) limit the possibility of identifying small bodies of liquid water or the existence of any hydraulic connection between them. Because of this, there is no reason to conclude that the presence of subsurface water on Mars is limited to a single location.

REFERENCES AND NOTES

1. S. M. Clifford, Polar basal melting on Mars. *J. Geophys. Res.* **92**, 9135–9152 (1987). doi:10.1029/JB092iB09p09135
2. D. W. Ashmore, R. G. Bingham, Antarctic subglacial hydrology: Current knowledge and future challenges. *Antarct. Sci.* **26**, 758–773 (2014). doi:10.1017/S0954102014000546
3. S. P. Carter, D. D. Blankenship, M. E. Peters, D. A. Young, J. W. Holt, D. L. Morse, Radar-based subglacial lake classification in Antarctica. *Geochem. Geophys. Geosyst.* **8**, 3016 (2007). doi:10.1029/2006GC001408
4. G. Picardi, J. J. Plaut, D. Biccari, O. Bombaci, D. Calabrese, M. Cartacci, A. Cicchetti, S. M. Clifford, P. Edenhofer, W. M. Farrell, C. Federico, A. Frigeri, D. A. Gurnett, T. Hagfors, E. Heggy, A. Herique, R. L. Huff, A. B. Ivanov, W. T. Johnson, R. L. Jordan, D. L. Kirchner, W. Kofman, C. J. Leuschen, E. Nielsen, R. Orosei, E. Pettinelli, R. J. Phillips, D. Plettemeier, A. Safaeinili, R. Seu, E. R. Stofan, G. Vannaroni, T. R. Watters, E. Zampolini, Radar soundings of the subsurface of Mars. *Science* **310**, 1925–1928 (2005). doi:10.1126/science.1122165 Medline
5. Materials and methods are available as supplementary materials.
6. W. M. Farrell, J. J. Plaut, S. A. Cummer, D. A. Gurnett, G. Picardi, T. R. Watters, A. Safaeinili, Is the Martian water table hidden from radar view? *Geophys. Res. Lett.* **36**, L15206 (2009). doi:10.1029/2009GL038945
7. J. J. Plaut, G. Picardi, A. Safaeinili, A. B. Ivanov, S. M. Milkovich, A. Cicchetti, W. Kofman, J. Mouginot, W. M. Farrell, R. J. Phillips, S. M. Clifford, A. Frigeri, R. Orosei, C. Federico, I. P. Williams, D. A. Gurnett, E. Nielsen, T. Hagfors, E. Heggy, E. R. Stofan, D. Plettemeier, T. R. Watters, C. J. Leuschen, P. Edenhofer, Subsurface radar sounding of the south polar layered deposits of Mars. *Science* **316**, 92–95 (2007). doi:10.1126/science.1139672 Medline
8. M. Cartacci, A. Frigeri, R. Orosei, E. Pettinelli, paper presented at the American Geophysical Union Fall Meeting, San Francisco, CA, 15 to 19 December 2008.
9. G. K. A. Oswald, S. P. Gogineni, Recovery of subglacial water extent from Greenland radar survey data. *J. Glaciol.* **54**, 94–106 (2008). doi:10.3189/002214308784409107
10. M. J. Siegert, S. Carter, I. Tabacco, S. Popov, D. D. Blankenship, A revised inventory of Antarctic subglacial lakes. *Antarct. Sci.* **17**, 453–460 (2005). doi:10.1017/S0954102005002889
11. D. E. Smith, M. T. Zuber, H. V. Frey, J. B. Garvin, J. W. Head, D. O. Muhleman, G. H. Pettengill, R. J. Phillips, S. C. Solomon, H. J. Zwally, W. B. Banerdt, T. C. Duxbury, M. P. Golombek, F. G. Lemoine, G. A. Neumann, D. D. Rowlands, O. Aharonson, P. G. Ford, A. B. Ivanov, C. L. Johnson, P. J. McGovern, J. B. Abshire, R. S. Afzal, X. Sun, Mars Orbiter Laser Altimeter: Experiment summary after the first year of global mapping of Mars. *J. Geophys. Res.* **106**, 23689–23722 (2001). doi:10.1029/2000JF001364
12. G. A. Neumann, J. B. Abshire, O. Aharonson, J. B. Garvin, X. Sun, M. T. Zuber, Mars Orbiter Laser Altimeter pulse width measurements and footprint-scale roughness. *Geophys. Res. Lett.* **30**, 1561 (2003). doi:10.1029/2003GL017048
13. C. S. Edwards, K. J. Nowicki, P. R. Christensen, J. Hill, N. Gorelick, K. Murray, Mosaicking of global planetary image datasets: 1. Techniques and data processing for Thermal Emission Imaging System (THEMIS) multi-spectral data. *J. Geophys. Res. Planets* **116**, E10008 (2011). doi:10.1029/2010JF003755
14. M. T. Zuber, R. J. Phillips, J. C. Andrews-Hanna, S. W. Asmar, A. S. Konopliv, F. G. Lemoine, J. J. Plaut, D. E. Smith, S. E. Smrekar, Density of Mars' south polar layered deposits. *Science* **317**, 1718–1719 (2007). doi:10.1126/science.1146995 Medline
15. J. Li, J. C. Andrews-Hanna, Y. Sun, R. J. Phillips, J. J. Plaut, M. T. Zuber, Density variations within the south polar layered deposits of Mars. *J. Geophys. Res. Planets* **117**, E04006 (2012).
16. M. L. Litvak, I. G. Mitrofanov, A. S. Kozyrev, A. B. Sanin, V. I. Tretyakov, W. V. Boynton, N. J. Kelly, D. Hamara, R. S. Saunders, Long-term observations of southern winters on Mars: Estimations of column thickness, mass, and volume density of the seasonal CO₂ deposit from HEND/Odyssey data. *J. Geophys. Res. Planets* **112**, E03S13 (2007). doi:10.1029/2006JF002832
17. O. Aharonson, M. Zuber, D. E. Smith, G. A. Neumann, W. C. Feldman, T. H. Prettyman, Depth, distribution, and density of CO₂ deposition on Mars. *J. Geophys. Res. Planets* **109**, E05004 (2004). doi:10.1029/2003JF002223
18. R. Seu, R. J. Phillips, D. Biccari, R. Orosei, A. Masdea, G. Picardi, A. Safaeinili, B. A. Campbell, J. J. Plaut, L. Marinangeli, S. E. Smrekar, D. C. Nunes, SHARAD sounding radar on the Mars Reconnaissance Orbiter. *J. Geophys. Res. Planets* **112**, E05S05 (2007). doi:10.1029/2006JF002745
19. R. Seu, R. J. Phillips, G. Alberti, D. Biccari, F. Bonaventura, M. Bortone, D. Calabrese, B. A. Campbell, M. Cartacci, L. M. Carter, C. Catallo, A. Croce, R. Croci, M. Cutigni, A. Di Placido, S. Dinardo, C. Federico, E. Flamini, F. Fois, A. Frigeri, O. Fuga, E. Giacomoni, Y. Gim, M. Guelfi, J. W. Holt, W. Kofman, C. J. Leuschen, L. Marinangeli, P. Marras, A. Masdea, S. Mattei, R. Mecozzi, S. M. Milkovich, A. Morlupi, J. Mouginot, R. Orosei, C. Papa, T. Paternò, P. Persi del Marmo, E. Pettinelli, G. Pica, G. Picardi, J. J. Plaut, M. Provenzani, N. E. Putzig, F. Russo, A. Safaeinili, G. Salzillo, M. R. Santovito, S. E. Smrekar, B. Tattarletti, D. Vicari, Accumulation and erosion of Mars' south polar layered deposits. *Science* **317**, 1715–1718 (2007). doi:10.1126/science.1144120 Medline
20. E. Pettinelli, B. Cosciotti, F. Di Paolo, S. E. Lauro, E. Mattei, R. Orosei, G. Vannaroni, Dielectric properties of Jovian satellite ice analogs for subsurface radar exploration: A review. *Rev. Geophys.* **53**, 593–641 (2015). doi:10.1002/2014RG000463
21. S. E. Lauro, E. Mattei, E. Pettinelli, F. Soldovieri, R. Orosei, M. Cartacci, A. Cicchetti, R. Noschese, S. Giuppi, Permittivity estimation of layers beneath the northern polar layered deposits, Mars. *Geophys. Res. Lett.* **37**, L14201 (2010). doi:10.1029/2010GL043015
22. M. A. Wiczorek, Constraints on the composition of the martian south polar cap

- from gravity and topography. *Icarus* **196**, 506–517 (2008). [doi:10.1016/j.icarus.2007.10.026](https://doi.org/10.1016/j.icarus.2007.10.026)
23. Z. Zhang, T. Hagfors, E. Nielsen, G. Picardi, A. Mesdea, J. J. Plaut, Dielectric properties of the Martian south polar layered deposits: MARSIS data inversion using Bayesian inference and genetic algorithm. *J. Geophys. Res. Planets* **113**, E05004 (2008). [doi:10.1029/2007JE002941](https://doi.org/10.1029/2007JE002941)
 24. L. M. Carter, B. A. Campbell, J. W. Holt, R. J. Phillips, N. E. Putzig, S. Mattei, R. Seu, C. H. Okubo, A. F. Egan, Dielectric properties of lava flows west of Ascraeus Mons, Mars. *Geophys. Res. Lett.* **36**, L23204 (2009). [doi:10.1029/2009GL041234](https://doi.org/10.1029/2009GL041234)
 25. J. Mouginot, A. Pommerol, W. Kofman, P. Beck, B. Schmitt, A. Herique, C. Grima, A. Safaeinili, J. J. Plaut, The 3–5 MHz global reflectivity map of Mars by MARSIS/Mars Express: Implications for the current inventory of subsurface H₂O. *Icarus* **210**, 612–625 (2010). [doi:10.1016/j.icarus.2010.07.003](https://doi.org/10.1016/j.icarus.2010.07.003)
 26. Y. Guéguen, V. Palciauskas, *Introduction to the Physics of Rocks* (Princeton Univ. Press, 1994).
 27. M. E. Peters, D. D. Blankenship, D. L. Morse, Analysis techniques for coherent airborne radar sounding: Application to West Antarctic ice streams. *J. Geophys. Res. Solid Earth* **110**, B06303 (2005). [doi:10.1029/2004JB003222](https://doi.org/10.1029/2004JB003222)
 28. A. A. Maryott, E. R. Smith, Table of Dielectric Constants of Pure Liquids (no. NBS-514, National Bureau of Standards, 1951).
 29. M. H. Hecht, S. P. Kounaves, R. C. Quinn, S. J. West, S. M. M. Young, D. W. Ming, D. C. Catling, B. C. Clark, W. V. Boynton, J. Hoffman, L. P. Deflores, K. Gospodinova, J. Kapit, P. H. Smith, Detection of perchlorate and the soluble chemistry of martian soil at the Phoenix lander site. *Science* **325**, 64–67 (2009). [doi:10.1126/science.1172466](https://doi.org/10.1126/science.1172466) [Medline](#)
 30. D. C. Catling, M. W. Claire, K. J. Zahnle, R. C. Quinn, B. C. Clark, M. H. Hecht, S. Kounaves, Atmospheric origins of perchlorate on Mars and in the Atacama. *J. Geophys. Res. Planets* **115**, E00E11 (2010). [doi:10.1029/2009JE003425](https://doi.org/10.1029/2009JE003425)
 31. Y. S. Kim, K. P. Wo, S. Maity, S. K. Atreya, R. I. Kaiser, Radiation-induced formation of chlorine oxides and their potential role in the origin of Martian perchlorates. *J. Am. Chem. Soc.* **135**, 4910–4913 (2013). [doi:10.1021/ja3122922](https://doi.org/10.1021/ja3122922) [Medline](#)
 32. D. A. Fisher, M. H. Hecht, S. P. Kounaves, D. C. Catling, A perchlorate brine lubricated deformable bed facilitating flow of the north polar cap of Mars: Possible mechanism for water table recharging. *J. Geophys. Res. Planets* **115**, E00E12 (2010). [doi:10.1029/2009JE003405](https://doi.org/10.1029/2009JE003405)
 33. J. Lasue, N. Mangold, E. Hauber, S. Clifford, W. Feldman, O. Gasnault, C. Grima, S. Maurice, O. Mousis, Quantitative assessments of the martian hydrosphere. *Space Sci. Rev.* **174**, 155–212 (2013). [doi:10.1007/s11214-012-9946-5](https://doi.org/10.1007/s11214-012-9946-5)
 34. C. Grima, F. Costard, W. Kofman, B. Saint-Bézar, A. Servain, F. Rémy, J. Mouginot, A. Herique, R. Seu, Large asymmetric polar scarps on Planum Australe, Mars: Characterization and evolution. *Icarus* **212**, 96–109 (2011). [doi:10.1016/j.icarus.2010.12.017](https://doi.org/10.1016/j.icarus.2010.12.017)
 35. R. Orosei, A. Cicchetti, Data files and electromagnetic simulation software used in the paper “Radar evidence of subglacial liquid water on Mars”. Zenodo (2018); <http://doi.org/10.5281/zenodo.1285179>.
 36. R. Jordan, G. Picardi, J. Plaut, K. Wheeler, D. Kirchner, A. Safaeinili, W. Johnson, R. Seu, D. Calabrese, E. Zampolini, A. Cicchetti, R. Huff, D. Gurnett, A. Ivanov, W. Kofman, R. Orosei, T. Thompson, P. Edenhofer, O. Bombaci, The Mars express MARSIS sounder instrument. *Planet. Space Sci.* **57**, 1975–1986 (2009). [doi:10.1016/j.pss.2009.09.016](https://doi.org/10.1016/j.pss.2009.09.016)
 37. G. Picardi, D. Biccari, R. Seu, J. Plaut, W. T. K. Johnson, R. L. Jordan, A. Safaeinili, D. A. Gurnett, R. Huff, R. Orosei, O. Bombaci, D. Calabrese, E. Zampolini, “MARSIS: Mars advanced radar for subsurface and ionosphere sounding,” in *Mars Express: The Scientific Payload*, A. Wilson, A. Chicarro, Eds. (ESA SP-1240, European Space Agency, 2004).
 38. A. Cicchetti, C. Nenna, J. J. Plaut, D. Plettemeier, R. Noschese, M. Cartacci, R. Orosei, Observations of Phobos by the Mars Express radar MARSIS: Description of the detection techniques and preliminary results. *Adv. Space Res.* **60**, 2289–2302 (2017). [doi:10.1016/j.asr.2017.08.013](https://doi.org/10.1016/j.asr.2017.08.013)
 39. M. Cartacci, E. Amata, A. Cicchetti, R. Noschese, S. Giuppi, B. Langlais, A. Frigeri, R. Orosei, G. Picardi, Mars ionosphere total electron content analysis from MARSIS subsurface data. *Icarus* **223**, 423–437 (2013). [doi:10.1016/j.icarus.2012.12.011](https://doi.org/10.1016/j.icarus.2012.12.011)
 40. J. C. Curlander, R. N. McDonough, *Synthetic Aperture Radar: Systems and Signal Processing* (Wiley, 1991).
 41. D. LeVine, The radar cross section of dielectric disks. *IEEE Trans. Antenn. Propag.* **32**, 6–12 (1984). [doi:10.1109/TAP.1984.1143180](https://doi.org/10.1109/TAP.1984.1143180)
 42. J. Mouginot, W. Kofman, A. Safaeinili, C. Grima, A. Herique, J. J. Plaut, MARSIS surface reflectivity of the south residual cap of Mars. *Icarus* **201**, 454–459 (2009). [doi:10.1016/j.icarus.2009.01.009](https://doi.org/10.1016/j.icarus.2009.01.009)
 43. A. Safaeinili, W. Kofman, J.-F. Nouvel, A. Herique, R. L. Jordan, Impact of Mars ionosphere on orbital radar sounder operation and data processing. *Planet. Space Sci.* **51**, 505–515 (2003). [doi:10.1016/S0032-0633\(03\)00048-5](https://doi.org/10.1016/S0032-0633(03)00048-5)
 44. D. C. Nunes, R. J. Phillips, Radar subsurface mapping of the polar layered deposits on Mars. *J. Geophys. Res. Planets* **111**, E06S21 (2006). [doi:10.1029/2005JE002609](https://doi.org/10.1029/2005JE002609)
 45. P. Lombardo, A. M. Magenta, E. Pettinelli, Multichannel fusion of subsurface radar images at different resolutions. *IET Radar Sonar Navig.* **147**, 121–133 (2000). [doi:10.1049/ip-rsn:20000196](https://doi.org/10.1049/ip-rsn:20000196)
 46. F. T. Ulaby, R. K. Moore, A. K. Fung, *Microwave Remote Sensing Fundamentals and Radiometry*, vol. 1 of *Microwave Remote Sensing: Active and Passive* (Addison-Wesley, 1981).
 47. C. Mätzler, “Microwave properties of ice and snow,” in *Solar System Ices*, B. Schmitt, C. de Bergh, M. Festou, Eds. (Springer, 1998), pp. 241–257.
 48. E. Pettinelli, G. Vannaroni, A. Cereti, F. Paolucci, G. Della Monica, M. Storini, F. Bella, Frequency and time domain permittivity measurements on solid CO₂ and solid CO₂-soil mixtures as Martian soil simulants. *J. Geophys. Res. Planets* **108**, 8029 (2003). [doi:10.1029/2002JE001869](https://doi.org/10.1029/2002JE001869)
 49. J. C. Maxwell Garnett, Colours in metal glasses and metal films. *Philos. Trans. R. Soc. Lond.* **203**, 385–420 (1904). [doi:10.1098/rsta.1904.0024](https://doi.org/10.1098/rsta.1904.0024)
 50. F. T. Ulaby, R. K. Moore, A. K. Fung, *From Theory to Applications*, vol. 3 of *Microwave Remote Sensing: Active and Passive* (Addison-Wesley, 1986).
 51. R. J. Phillips, B. J. Davis, K. L. Tanaka, S. Byrne, M. T. Mellon, N. E. Putzig, R. M. Haberle, M. A. Kahre, B. A. Campbell, L. M. Carter, I. B. Smith, J. W. Holt, S. E. Smrekar, D. C. Nunes, J. J. Plaut, A. F. Egan, T. N. Titus, R. Seu, Massive CO₂ ice deposits sequestered in the south polar layered deposits of Mars. *Science* **332**, 838–841 (2011). [doi:10.1126/science.1203091](https://doi.org/10.1126/science.1203091) [Medline](#)
 52. A. Dobrovolskis, A. P. Ingersoll, Carbon Dioxide-Water Clathrate as a Reservoir of CO₂ on Mars. *Icarus* **26**, 353–357 (1975). [doi:10.1016/0019-1035\(75\)90178-5](https://doi.org/10.1016/0019-1035(75)90178-5)
 53. J. L. Whitten, B. A. Campbell, Lateral continuity of layering in the Mars South Polar Layered Deposits from SHARAD sounding data. *J. Geophys. Res. Planets* **10.1029/2018JE005578** (2018).

ACKNOWLEDGMENTS

We gratefully acknowledge the work of Giovanni Picardi (1936–2015), who served as Principal Investigator of MARSIS. The MARSIS instrument and experiment were funded by the Italian Space Agency and NASA and developed by the University of Rome, Italy, in partnership with NASA’s Jet Propulsion Laboratory (JPL), Pasadena, CA. Alenia Spazio (now Thales Alenia Space, Italy) provided the instrument’s digital processing system and integrated the parts and now operates the instrument and experiment. The University of Iowa, Iowa City, IA, built the transmitter for the instrument; JPL built the receiver; and Astro Aerospace, Carpinteria, CA, built the antenna. This research has made use of NASA’s Astrophysics Data System. The perceptually uniform color map “broc” was used in this study to prevent visual distortion of the data. We thank M. Mastrogioseppe and G. Vannaroni for insightful discussions. We are grateful to S. E. Beaubien for careful proofreading of the manuscript and improvement of the English. **Funding:** This work was supported by the Italian Space Agency (ASI) through contract I/032/12/1. M.P. acknowledges the support from the NASA Postdoctoral Program (2015–2017) at the Ames Research Center in Moffett Field, California. **Author contributions:** R.O. devised the data calibration method, produced maps of subsurface reflectors, developed the electromagnetic propagation model, codeveloped the method for data interpretation, and cowrote the paper. S.E.L. contributed to the development of the electromagnetic propagation model, codeveloped the method for data interpretation, and cowrote the paper. E.P. coordinated the writing of the paper, contributed to data analysis interpretation, and discussed ideas. A.C. planned and conducted the search for bright subsurface radar reflectors using raw data. M.Co., B.C., F.D.P., E.F., E.M., and M.P. contributed text and figures to the manuscript and discussed ideas. F.S. contributed to the forward and inverse modeling of the electromagnetic propagation and scattering and discussed

ideas. M.Ca., F.C., A.F., S.G., R.M., A.M., G.M., C.N., R.N., M.R., and R.S. contributed to data acquisition and analysis and discussed ideas. **Competing interests:** The authors declare no competing interests. **Data and materials availability:** Data reported in this paper, scripts used to model electromagnetic propagation, and the output of those scripts are available through the Zenodo research data repository (35).

SUPPLEMENTARY MATERIALS

www.sciencemag.org/cgi/content/full/science.aar7268/DC1

Materials and Methods

Supplementary Text

Figs. S1 to S6

Table S1

References (36–53)

13 December 2017; accepted 20 June 2018

Published online 25 July 2018

10.1126/science.aar7268

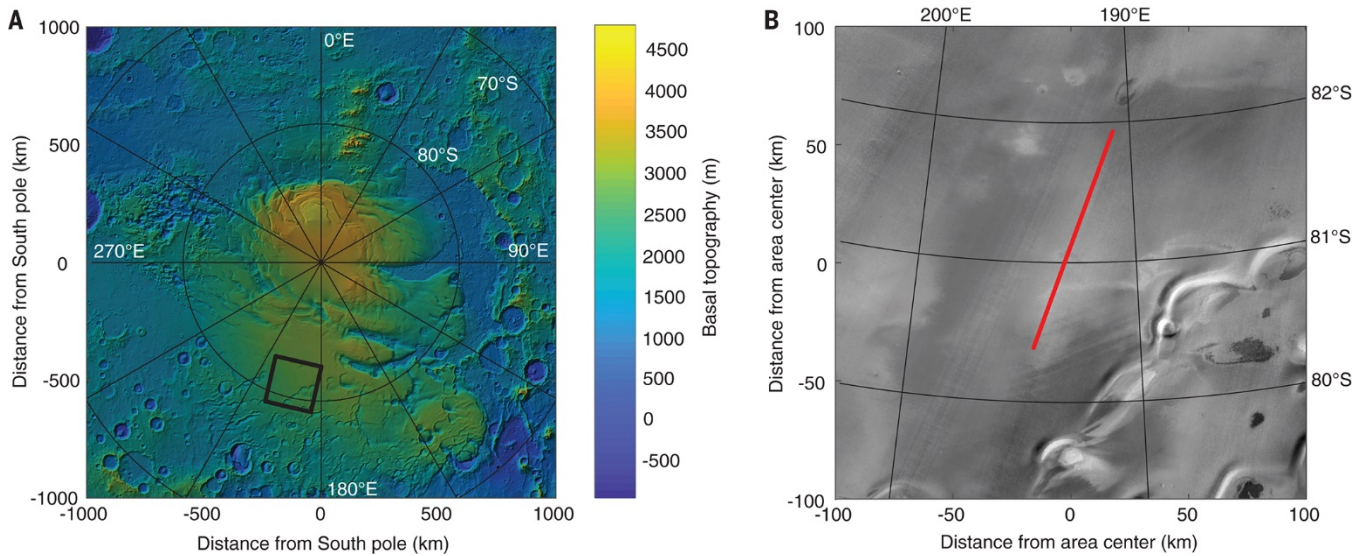


Fig. 1. Maps of the investigated area. (A) Shaded relief map of Planum Australe, Mars, south of 75°S latitude. The map was produced using the MOLA topographic dataset (11). The black square outlines the study area. (B) Mosaic produced using infrared observations by the THEMIS (Thermal Emission Imaging System) camera (13), corresponding to the black square in (A). South is up in the image. The red line marks the ground track of orbit 10737, corresponding to the radargram shown in Fig. 2A. The area consists mostly of featureless plains, except for a few large asymmetric polar scarps near the bottom right of (B), which suggest an outward sliding of the polar deposits (34).

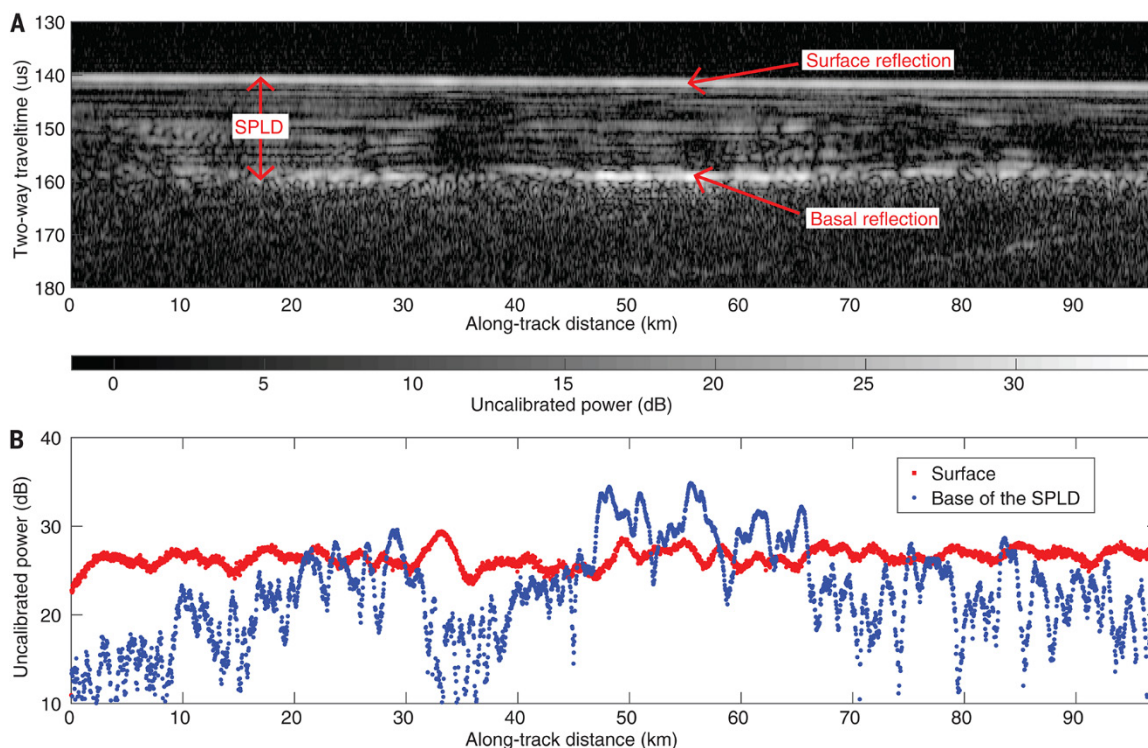


Fig. 2. Radar data collected by MARSIS. (A) Radargram for MARSIS orbit 10737, whose ground track is shown in Fig. 1B. A radargram is a bi-dimensional color-coded section made of a sequence of echoes in which the horizontal axis is the distance along the ground track of the spacecraft, the vertical axis represents the two-way travel time of the echo (from a reference altitude of 25 km above the reference datum), and brightness is a function of echo power. The continuous bright line in the topmost part of the radargram is the echo from the surface interface, whereas the bottom reflector at about 160 μ s corresponds to the SPLD/basal material interface. Strong basal reflections can be seen at some locations, where the basal interface is also planar and parallel to the surface. (B) Plot of surface and basal echo power for the radargram in (A). Red dots, surface echo power; blue dots, subsurface echo power. The horizontal scale is along-track distance, as in (A), and the vertical scale is uncalibrated power in decibels. The basal echo between 45 and 65 km along-track is stronger than the surface echo even after attenuation within the SPLD.

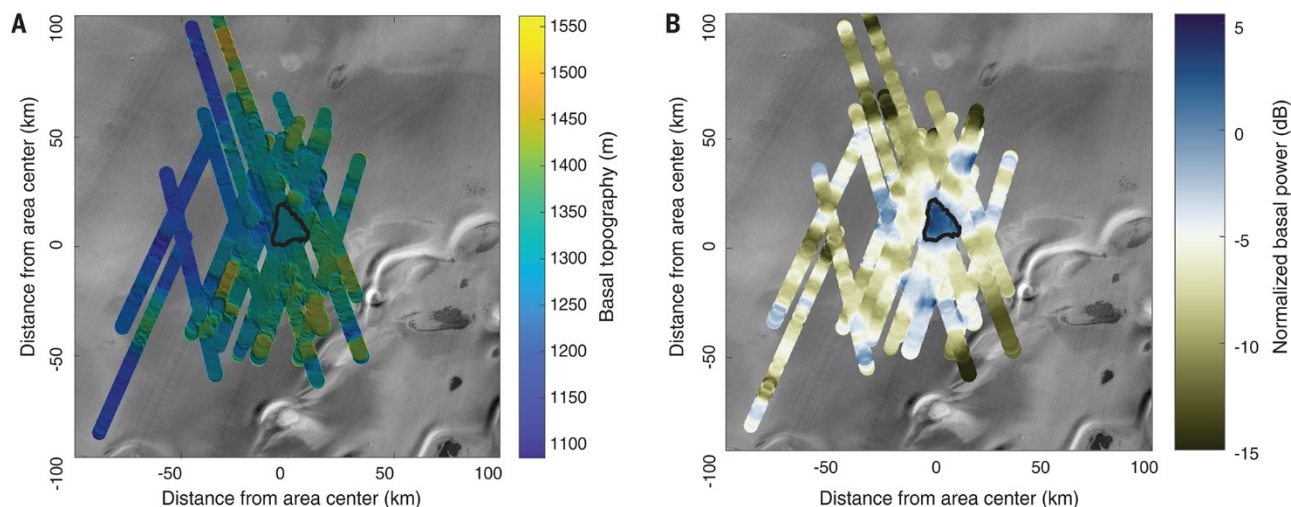


Fig. 3. Maps of basal topography and reflected echo power. (A) Color-coded map of the topography at the base of the SPLD, computed with respect to the reference datum. The black contour outlines the area in which bright basal reflections are concentrated. (B) Color-coded map of normalized basal echo power at 4 MHz. The large blue area (positive values of the normalized basal echo power) outlined in black corresponds to the main bright area; the map also shows other, smaller bright spots that have a limited number of overlapping profiles. Both panels are superimposed on the infrared image shown in Fig. 1B, and the value at each point is the median of all radar footprints crossing that point.

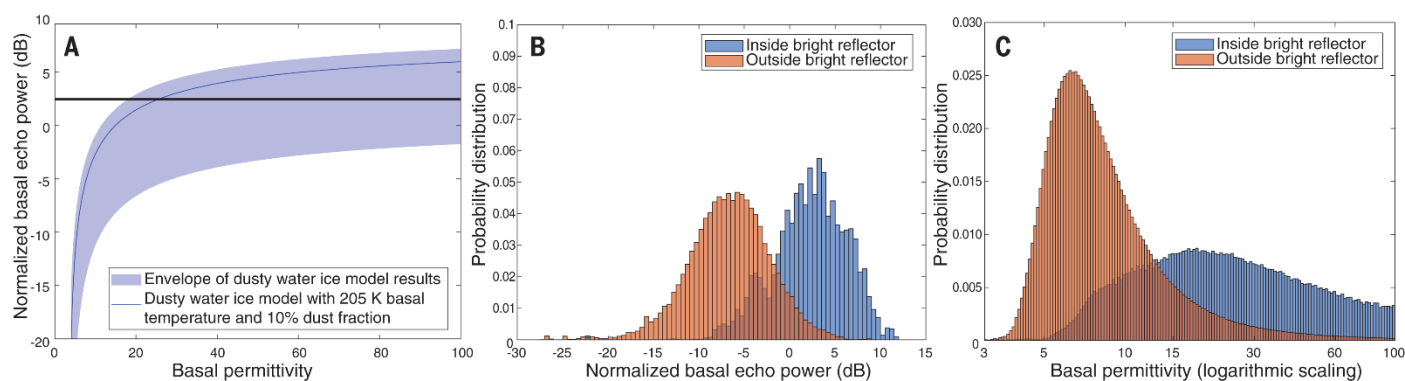


Fig. 4. Results of the simulation and retrieved permittivities. (A) Output of the electromagnetic simulations computed at 4 MHz (figs. S4 and S6). The blue shaded area is the envelope of all curves incorporating different amounts of H_2O ice and dust along with various basal temperatures for the SPLD. The blue line is the curve for a single model (basal temperature of 205 K and 10% dust content), shown for illustration, and the black horizontal line is the median normalized basal echo power at 4 MHz from the observations. (B) Normalized basal echo power distributions inside (blue) and outside (brown) the bright reflection area, indicating two distinct populations of values. These distributions, together with the chart in (A), are used to compute the basal permittivity; for example, the intersection between the blue curve and the black line gives a basal permittivity value of 24. (C) Basal permittivity distributions inside (blue) and outside (brown) the bright reflection area. The nonlinear relationship between the normalized basal echo power and the permittivity produces an asymmetry (skewness) in the distributions of the values.

Radar evidence of subglacial liquid water on Mars

R. Orosei, S. E. Lauro, E. Pettinelli, A. Cicchetti, M. Coradini, B. Cosciotti, F. Di Paolo, E. Flamini, E. Mattei, M. Pajola, F. Soldovieri, M. Cartacci, F. Cassenti, A. Frigeri, S. Giuppi, R. Martufi, A. Masdea, G. Mitri, C. Nenna, R. Noschese, M. Restano and R. Seu

published online July 25, 2018

ARTICLE TOOLS	http://science.sciencemag.org/content/early/2018/07/24/science.aar7268
SUPPLEMENTARY MATERIALS	http://science.sciencemag.org/content/suppl/2018/07/24/science.aar7268.DC1
RELATED CONTENT	http://science.sciencemag.org/content/sci/361/6400/320.full http://science.sciencemag.org/content/sci/361/6401/448.full
REFERENCES	This article cites 42 articles, 6 of which you can access for free http://science.sciencemag.org/content/early/2018/07/24/science.aar7268#BIBL
PERMISSIONS	http://www.sciencemag.org/help/reprints-and-permissions

Use of this article is subject to the [Terms of Service](#)